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# EISCAT measurements of the solar wind

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**Abstract.** EISCAT observations of interplanetary scintillation have been used to measure the velocity of the solar wind at distances between 15 and 130  $R_{\odot}$  (solar radii) from the Sun. The results show that the solar wind consists of two distinct components, a fast stream with a velocity of  $\sim 800 \text{ km s}^{-1}$  and a slow stream at  $\sim 400 \text{ km s}^{-1}$ . The fast stream appears to reach its final velocity much closer to the Sun than expected. The results presented here suggest that this is also true for the slow solar wind. Away from interaction regions the flow vector of the solar wind is purely radial to the Sun. Observations have been made of fast wind/slow wind interactions which show enhanced levels of scintillation in compression regions.

## 1 Observations

Measurements of IPS (interplanetary scintillation) have been used for many years to study electron density fluctuations in the solar wind (e.g. Cohen *et al.*, 1967) and to measure the solar wind velocity (Dennison and Hewish, 1967; Armstrong and Coles, 1972; Coles 1995). IPS has the advantage that it can provide measurements from close to the Sun to beyond 1 AU and at all heliographic latitudes (Coles, 1995). The solar wind velocity is determined from the correlation of scintillations observed at separate sites and the three EISCAT receiving stations provide an ideal system, offering baselines up to 390 km. The frequency band used for these observations is centred at 931.5 MHz, allowing measurements from as close to the Sun as 15  $R_{\odot}$  (solar radii) out to distances of more than 130  $R_{\odot}$ .

IPS observations were first made at EISCAT in 1981 (Bourgois *et al.*, 1985), and since 1990 a regular series of measurements has been made every summer (Coles *et al.*, 1991; Breen *et al.*, 1996; Grall *et al.*, 1996). The measurements made in 1994 and 1995 were especially significant

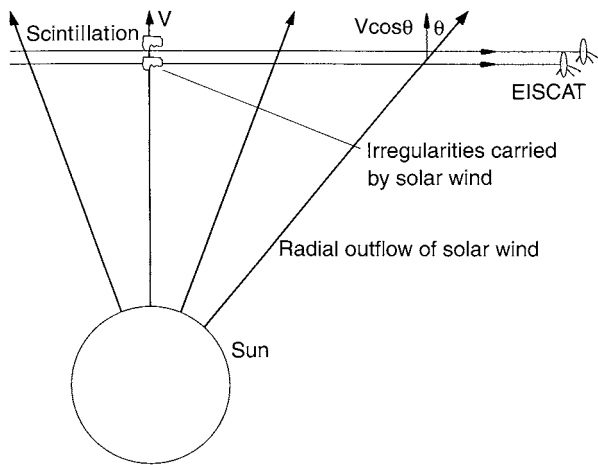
as they matched the passage of the ULYSSES spacecraft over the poles of the Sun.

The data used in this study were taken between April 1992 and September 1995 using the EISCAT special programme SP-UK-SCINTI to monitor the scintillations in signals from a radio source of small angular diameter received at two of the EISCAT UHF antennas. If the parallel lines of sight to the two antennas pass through the solar atmosphere in the same radial plane then the fluctuations in signal power received by the antennas are highly correlated. As the distance between the two antennas is known, the time lag for maximum cross-correlation can be used to estimate the component of solar wind velocity perpendicular to the line of sight (Coles *et al.*, 1991; Breen *et al.*, 1996). The long baselines available between EISCAT sites mean that these velocity measurements can be made with high resolution so that streams of different velocity can be distinguished (Grall *et al.*, 1996).

## 2 Results

### 2.1 Bimodal velocity distribution

In earlier IPS work using EISCAT (e.g. Bourgois *et al.*, 1985) the apparent velocity was determined by fitting the observed correlation functions with a theoretical model in which the solar wind velocity had a compact distribution about a single mean speed. The mean speed was largely determined from the time lag at which the auto- and cross-correlation functions intersect. The rms velocity in the tangential direction causes the peak of the cross-correlation function to drop, but does not increase its width. The rms velocity in the radial direction does not change the area under the cross-correlation function but broadens and skews it, causing the peak cross-correlation to drop. The rms velocity in the tangential direction could be caused by plasma waves, whereas the rms velocity in the radial direction is largely due to the 'cos  $\theta$ ' effect, where  $\theta$  is the angle between the local direction of the solar wind and the direction perpendicular to the line of sight from

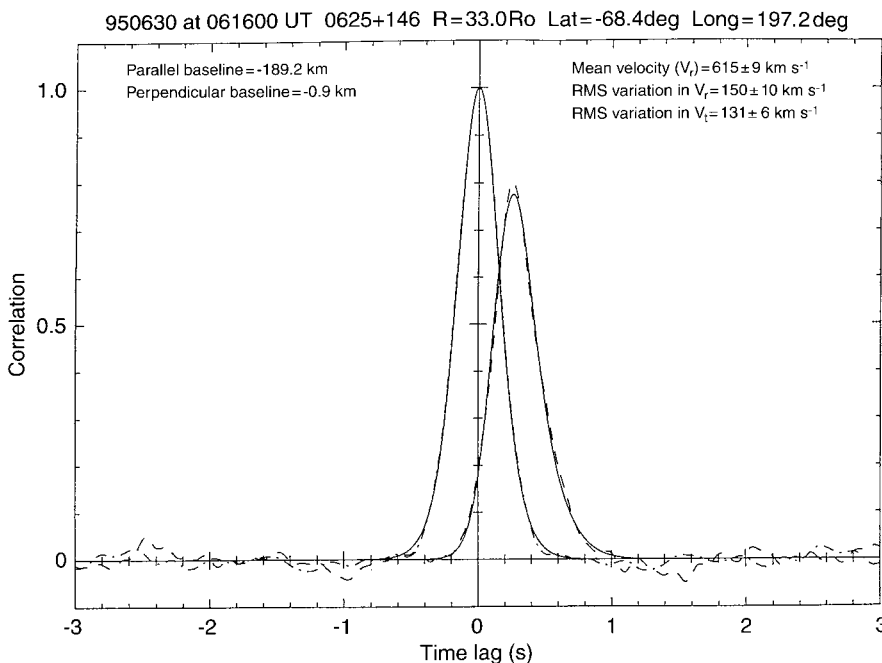


**Fig. 1.** Schematic diagram of EISCAT IPS observations (SP-UK-SCINTI). Density fluctuations carried outwards by the solar wind modulate the radio waves from the distant source, producing a moving interference pattern at the receiving sites. The experiment is only sensitive to the component of solar wind velocity perpendicular to the line of sight to the source ( $V \cos \theta$ , where  $V$  is the solar wind velocity)

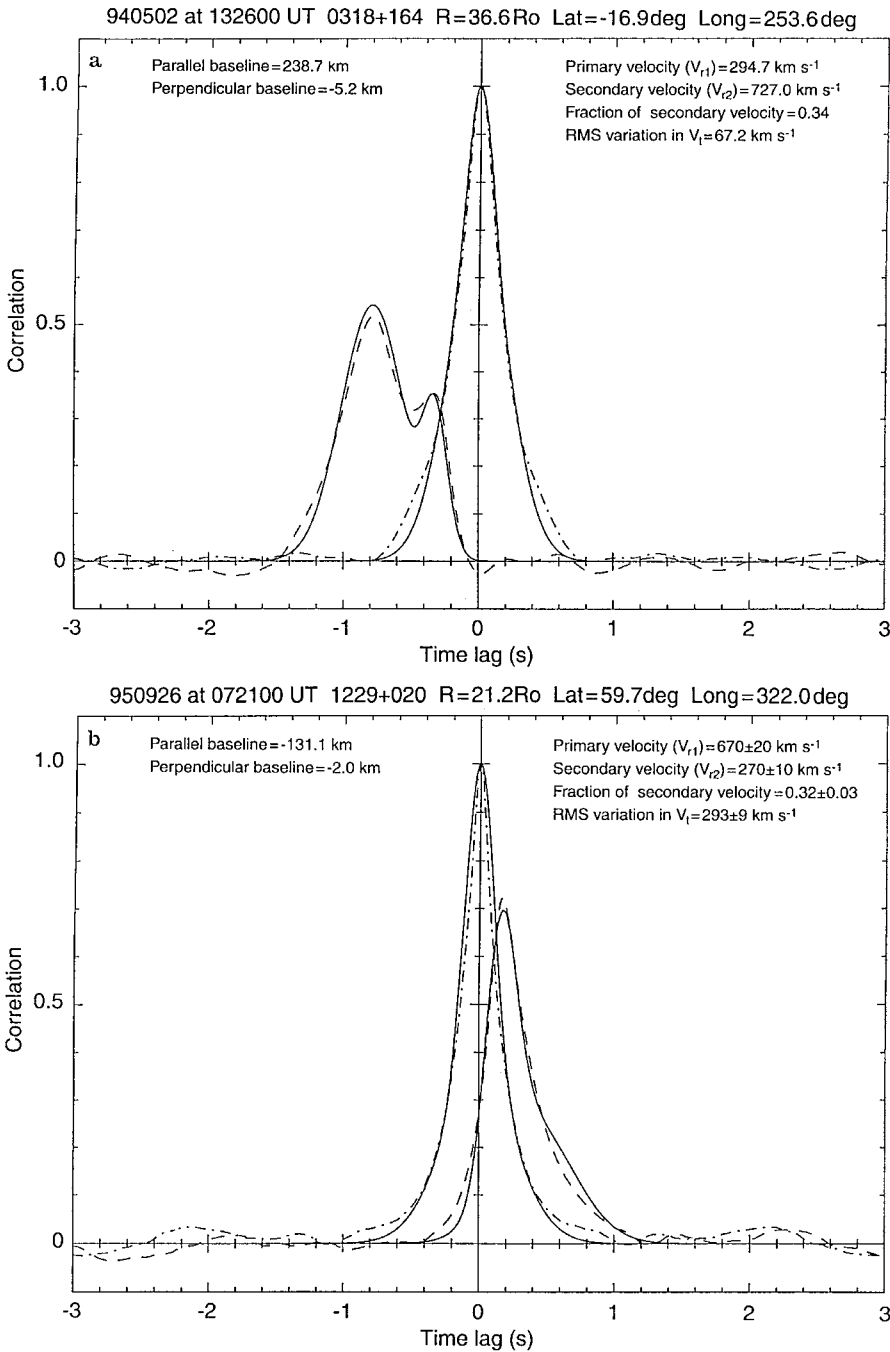
the source to the antenna (Fig. 1). Fortunately the strength of scattering falls rapidly with distance from the Sun ( $\propto 1/R^4$ ) so that in the single-velocity case the strongest scattering occurs for small values of  $\theta$  where the  $\cos \theta$  effect is small. The true radial velocity can then be derived from the apparent velocity by applying a correction factor of about 1.18. Under ideal conditions, the cross-correlation function for a single-velocity solar wind is narrow and almost symmetrical with a high peak value and in such cases both the rms deviations are small and  $V$  can be measured with a random error of only a few percent (see Fig. 2).

In contrast, other cases show a cross-correlation function which is rather wide, with a strong asymmetry. When such a function is analysed in terms of a single-velocity solar wind the results suggest rms deviations in both radial and transverse velocities that are much larger than appear realistic. However, from 1994 onwards the EISCAT observations were planned to match those made by Ulysses at high latitude and on the whole the baselines used were longer than in previous years. For observations made at such long baselines the cross-correlation functions often showed two distinct peaks, suggesting that two separate velocity streams were present along the line of sight (Breen *et al.*, 1996; Grall *et al.*, 1996). As a result a fitting procedure was adopted which assumed the presence of two different streams and this produced model correlation functions which fitted closely to the two-peak functions actually measured (Fig. 3a). In the light of this result the entire dataset was re-analysed assuming the presence of a fast stream and a slow stream in the solar wind, both showing small intrinsic variations in velocity. Overall the accuracy of the model fits was significantly improved when such a two-stream model was used, especially in the case of asymmetric functions like the one shown in Fig. 3b. In this case, the new analysis programme indicated that 62% of the scintillation was caused by a fast stream with an apparent velocity of  $670 \text{ km s}^{-1}$  and 38% from a slow stream with an apparent velocity of  $300 \text{ km s}^{-1}$ .

Figure 4a shows a frequency histogram of the different velocities derived by the new analysis programme for EISCAT measurements at low (below  $30^\circ$  north or south) and high (above  $60^\circ$ ) heliographic latitudes made between 1992 and 1995. Two clear peaks are present, one corresponding to low-latitude measurements with apparent velocities of about  $300 \text{ km s}^{-1}$  and the other to high-latitude measurements with apparent velocities of about



**Fig. 2.** Observations of 0625+146 on 30 June 1995, showing a clear fast stream at high latitude (single stream analysis). The dashed line represents the data and the solid line the model fit



**Fig. 3.** **a** Observations of 0318 + 164 on 2 May 1994, showing a two-peak correlation function analysed using the two-velocity model. The *dashed line* represents the data and the *solid line* the model fit; **b** Observations of 1229 + 020 on 26 September 1995, showing a skewed single peak correlation function analysed using the two-velocity model. The *dashed line* represents the data and the *solid line* the model fit

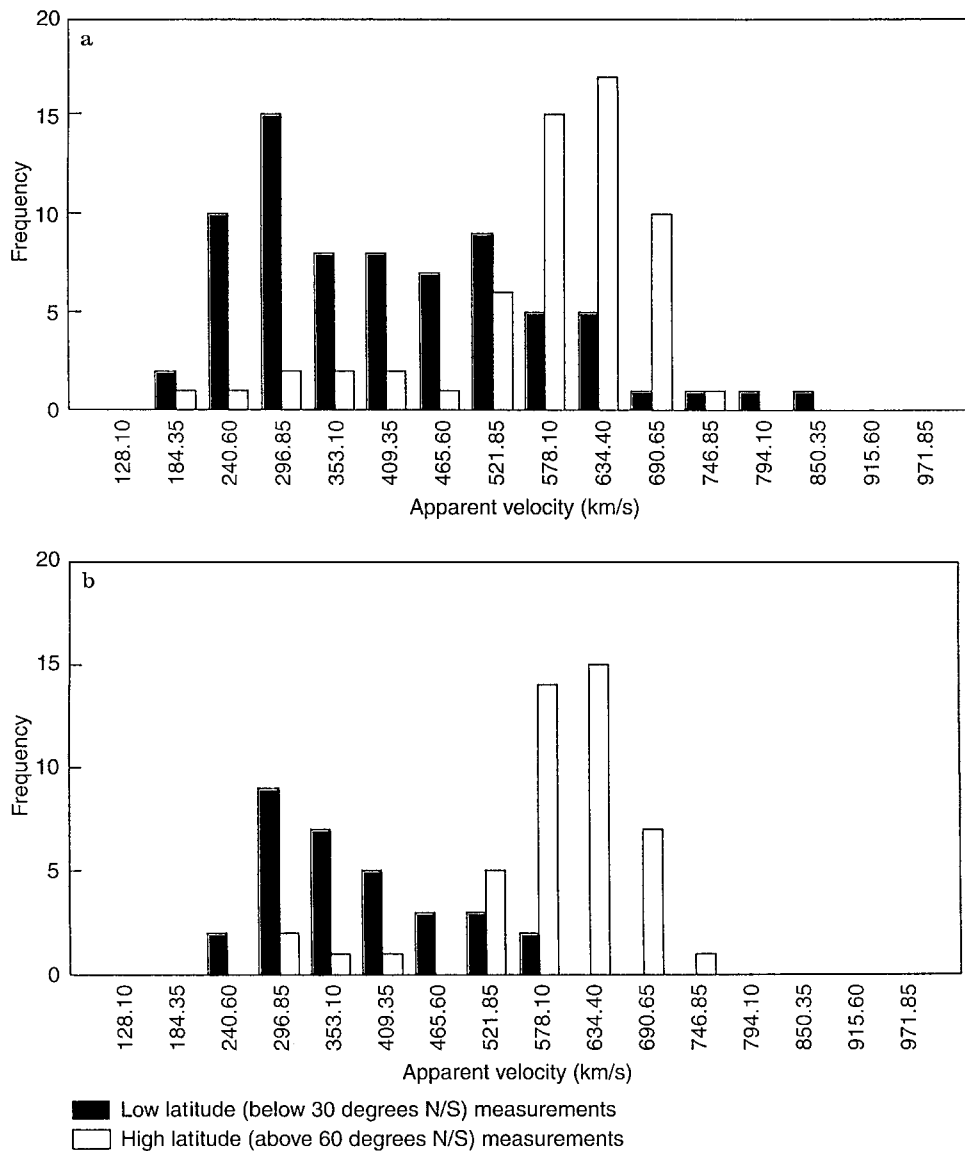
600 km s<sup>-1</sup>. These peaks become even more marked when only those cases are included where one stream is clearly dominant, contributing more than 70% of the observed scintillation power, as shown in Fig. 4b. The low-latitude observations are clearly dominated by slow streams, while the high-latitude observations are dominated by fast streams with apparent velocities of 600–700 km s<sup>-1</sup>. When corrected for foreshortening these measurements imply true velocities of 700–800 km s<sup>-1</sup>, which is consistent with Ulysses observations of fast streams at high heliographic latitudes (Phillips *et al.*, 1994).

These results indicate that the solar wind velocity falls into two categories with distinct velocities of about 400 and 800 km s<sup>-1</sup> rather than varying continuously over

a wide range of speeds. This bimodal distribution provides an interpretation of IPS data consistent with the description of a fast solar wind originating from coronal holes and a slow solar wind from the “non-hole” corona. (Sheeley *et al.*, 1976). It is also consistent with the observations made at much greater solar distances by spacecraft such as Helios (Schwenn, 1990) and Ulysses (Phillips *et al.*, 1994, 1995).

## 2.2 Variation of velocity with distance from Sun

The IPS technique has the great advantage of providing repeated measurements of solar wind velocity over a wide



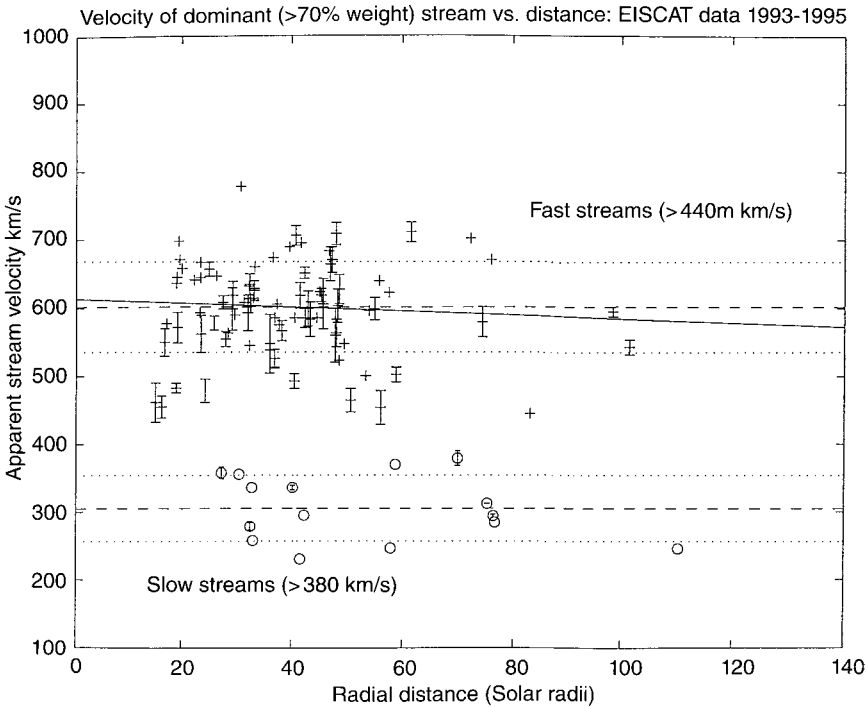
**Fig. 4 a, b.** Histograms of apparent solar wind velocity measurements for high ( $>60^\circ$  N/S) and low ( $<30^\circ$  N/S) heliographic latitudes, 1992–1995. **a** All measurements; **b** all cases where one velocity is dominant, contributing  $>70\%$  of the scattering

range of distances from the Sun. However, the fitting procedure assumes weak scattering along the line of sight, and at any specified frequency there is a minimum distance inside which strong scattering occurs so this procedure cannot be used. EISCAT operates at a relatively high frequency for an IPS system and so is able to make measurements relatively close to the Sun. In fast streams (for example, those originating at the solar poles) weak scattering can be assumed as close to the Sun as  $15 R_\odot$ . The slow (equatorial) streams are denser by a factor of  $\sim 3$  (Axford and McKenzie, 1996) so that the strength of scattering per unit volume is greater and the velocity of the slow solar wind can only be measured with certainty beyond  $30 R_\odot$ .

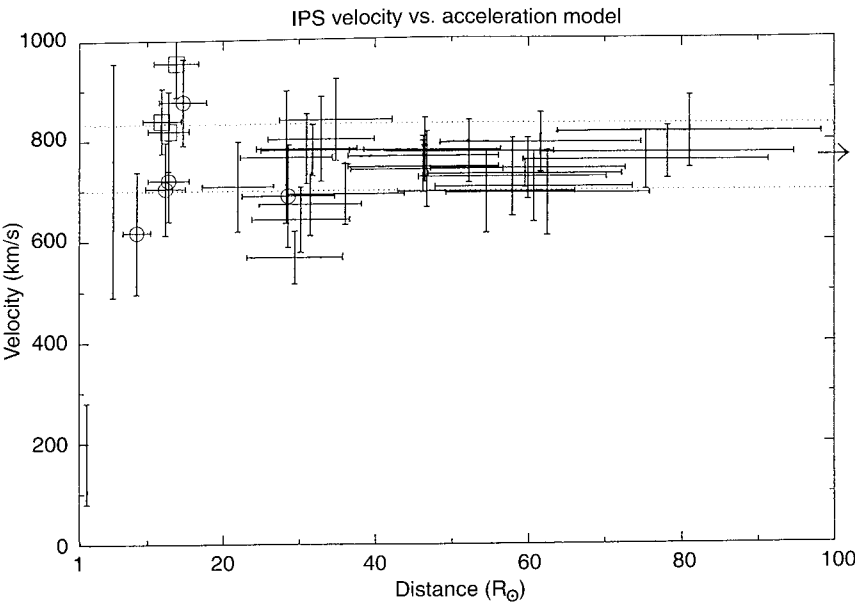
Figure 5 shows the variation of the apparent velocity of dominant fast and slow streams of solar wind with distance from the Sun. The broken horizontal lines represent the mean apparent velocity of the fast and slow streams. There is no evidence of any variation in velocity over the

range of distances observed. Furthermore, when corrected for foreshortening, the fast and slow stream velocities are in close agreement with those observed by Ulysses at distances of 1.5 AU or more (Phillips *et al.*, 1994). The error bars shown in Fig. 5 are based on the fit of the modelled correlation functions to the data. They do not take account of any other sources of error so the error bars shown in Fig. 5 represent the minimum uncertainties in the velocity estimates.

These results suggest that the solar wind completes its acceleration within a distance less than  $15\text{--}30 R_\odot$ , a conclusion strengthened by the results of a series of co-ordinated observations using EISCAT and the VLBA system in New Mexico. The VLBA operates at frequencies considerably higher than those available to EISCAT and hence can make reliable velocity measurements as close to the Sun as  $2.4 R_\odot$ . Figure 6 shows the results of a series of measurements of the high-latitude solar wind made using EISCAT and the VLBA (Grall *et al.*, 1996). The fast



**Fig. 5.** Variation with radial distance of the apparent velocity of clearly dominant streams (contributing >85% of scattering) in data from 1993–1995



**Fig. 6.** Variation of velocity of high-latitude fast streams determined from EISCAT and VLBA measurements (Grall *et al.*, 1996). The EISCAT measurements (from the southern polar hole between April and September 1994) are marked with *crossed bars*, while VLBA measurements are marked with *circles* (south polar hole measurements) or *squares* (measurements from the northern polar hole). In all cases the *vertical bar* represents the 90% confidence limit for the velocity and the *horizontal bar* indicates the distance range over which the measurement is averaged. The *points* inside  $10 R_{\odot}$  are derived from SPARTAN 201-01 white-light measurements (Fisher and Guhathakurta, 1994; Habbal *et al.*, 1995)

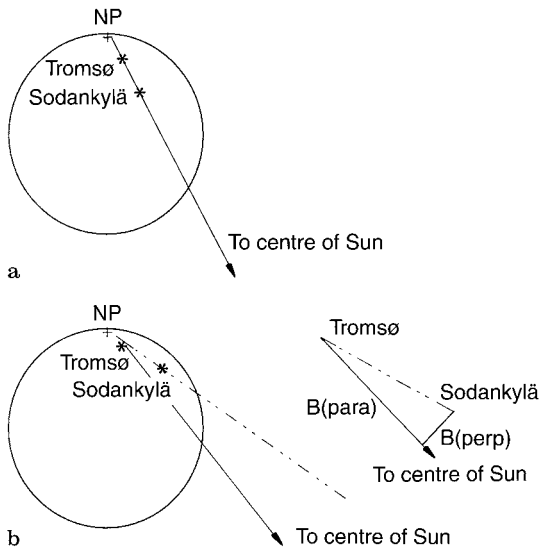
stream velocity (corrected for foreshortening) is plotted against distance from the Sun. An arrow at  $100 R_{\odot}$  denotes the mean velocity measured by Ulysses over the same latitude range.

### 2.3 Direction of flow

Theories of the solar wind suggest that away from interaction regions the solar wind should flow radially outwards from the Sun (e.g. Axford and McKenzie, 1996) so we would expect the maximum cross-correlation to

occur when the lines of sight from the source to the two antennas lie in a radial plane passing through the centre of the Sun, as shown in Fig. 7a. In the normal mode of operation EISCAT observes a radio source for 15 min, centred at the optimum time when the baseline lies in this radial plane. However, as the Earth rotates the two lines of sight are no longer in the same radial plane (Fig. 7b) and if the observations are extended for several hours the cross-correlation between the scintillations at the two sites will eventually disappear.

A number of these extended observations were carried out in 1994, and in each case the maximum correlation



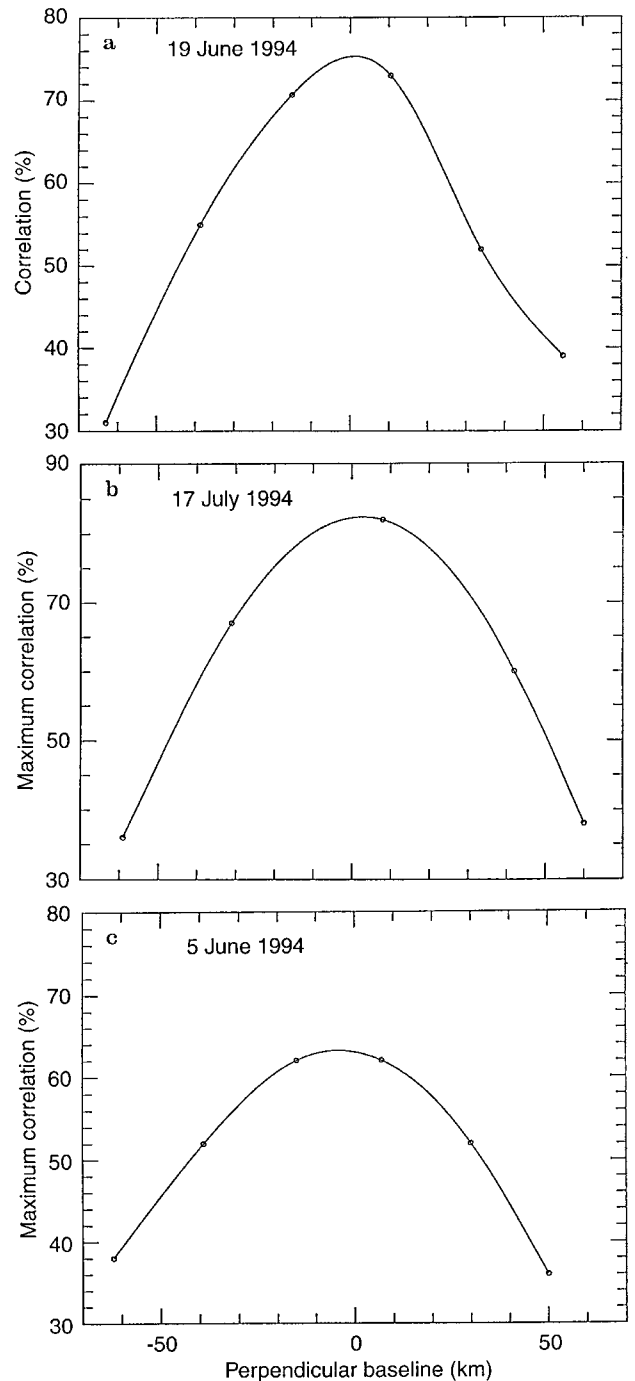
**Fig. 7.** **a** The EISCAT antennas viewed from the direction of the source. The transverse separation between the two sites (Tromsø and Sodankylä) is zero. **NP** marks the geographical north pole of the Earth. **b** The EISCAT antennas viewed from the direction of the source 2 h later than **a**. The transverse separation between the two sites is no longer zero

coefficient was plotted against the “perpendicular baseline” (the transverse displacement of one line of sight from the radial plane defined by the other,  $B(\text{perp})$  in the enlarged section of Fig. 7b). After correction for the velocity of the Earth and for the movement of the antennas during the time lag to maximum correlation, the highest correlation was found, in almost every case, to correspond to a zero perpendicular baseline, as shown in the examples given as Fig. 8a, b. On one occasion however a small offset was observed (Fig. 8c), suggesting that the solar wind had been deflected from the radial direction by  $1\text{--}1.5^\circ$ . This is a very small deflection but the direction is consistent with the effects of a fast stream/slow stream interaction.

#### 2.4 Association with coronal features

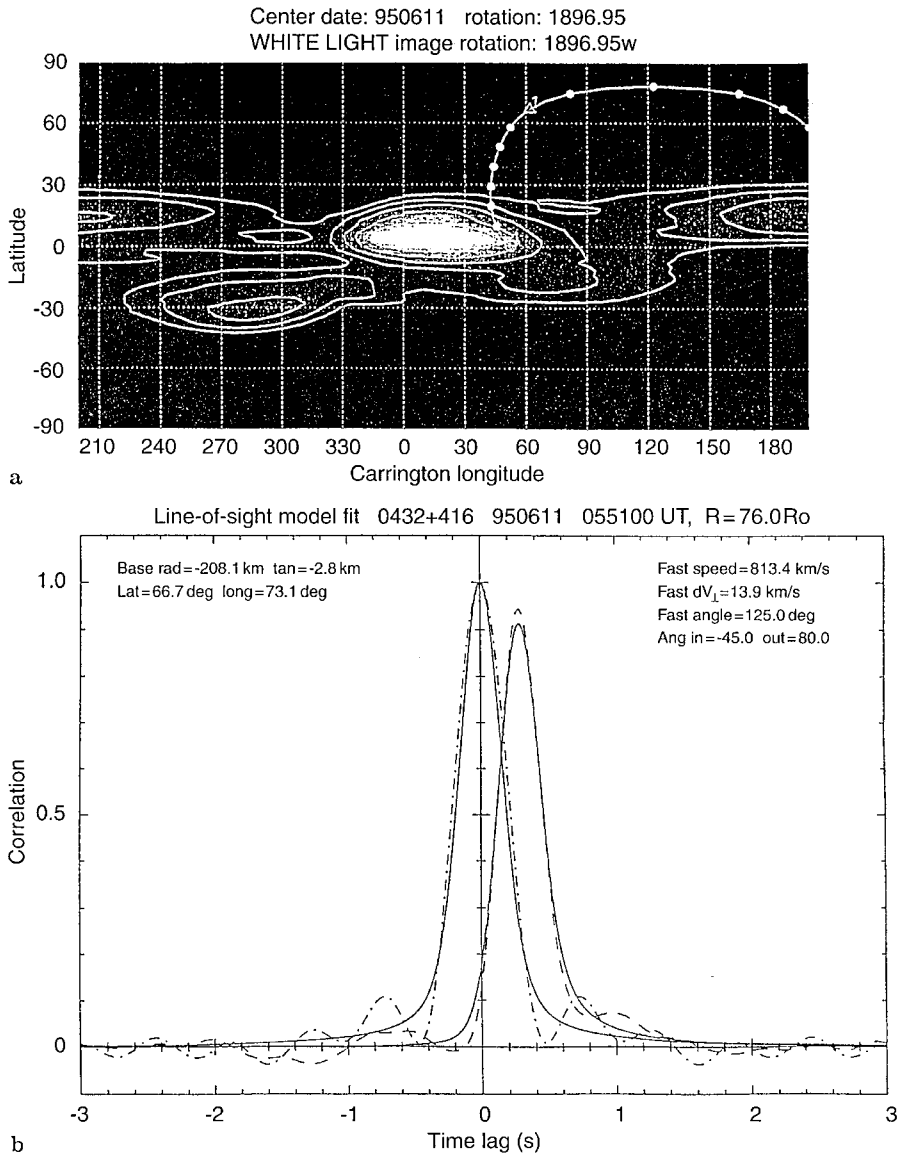
The results summarised indicate that the solar wind flows radially (or close to radially) with a velocity of  $800\text{ km s}^{-1}$  (fast stream) or  $400\text{ km s}^{-1}$  (slow stream), and that these velocities are approximately constant from inside  $15 R_\odot$  (fast streams) or  $25 R_\odot$  (slow streams). Knowing the rate of rotation of the Sun, these results make it practical to ‘map’ each line of sight from source to antenna down into the upper corona ( $2\text{--}5 R_\odot$ ). This in turn allows the region of origin for each part of the solar wind along the line of sight to be compared with white-light or X-ray observations of the corona. As IPS is essentially a limb observation the white-light coronagraph images produced by the High-Altitude Observatory are particularly suitable for comparison, although the full disk X-ray images produced by the Yohkoh spacecraft can also be used.

Figure 9 shows the white-light map of the corona produced from west-limb observations of the Sun made by the HAO. The central date of the observations is 11



**Fig. 8.** **a** Maximum cross-correlation versus perpendicular baseline for observations of 0625 + 146 on 19 June 1994. When the velocity of the Earth is taken into account the maximum cross-correlation between the measurements occurs when the perpendicular baseline is zero, indicating that the solar wind velocity is purely radial. **b** As **a** but showing observations of 0745 + 101 on 17 July 1994; **c** maximum cross-correlation versus perpendicular baseline for observations of 0521 + 166 on 5 June 1994. The maximum cross-correlation occurs for a small negative perpendicular baseline, suggesting that the solar wind velocity may be  $1.0^\circ$  to  $1.5^\circ$  offradial

June 1995. The map is plotted in a ‘natural’ grey-scale, with the bright equatorial region appearing in white and the large polar coronal holes as dark areas. Contour lines of constant brightness are added as a guide. The solid line



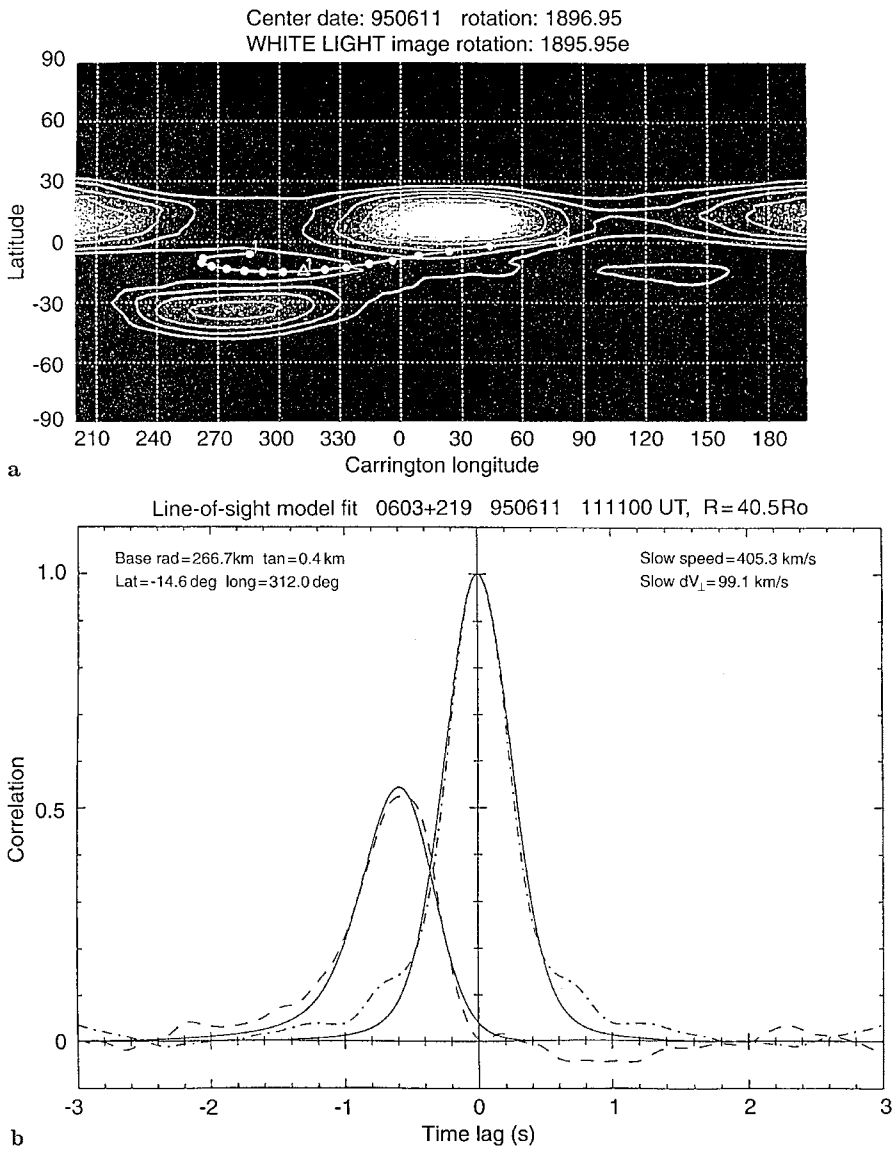
**Fig. 9.** **a** White-light map with the IPS line of sight for observations of 0432 + 416 on 11 June 1995 (a single fast stream) back-projected onto the map with an assumed constant radial velocity of  $800 \text{ km s}^{-1}$ . As in Fig. 8 the closest approach of the line of sight to the Sun is marked with a triangle and the position of the Earth is shown by  $\oplus$ . The line of sight enters the fast stream at  $-45^\circ$  and remains in it for the rest of its length. **b** IPS correlation function corresponding to **a**, analysed using the line of sight model. The dashed line represents the data and the solid line the model fit

superimposed over the map is the line of sight through the solar wind from the radio source 0432 + 416 to EISCAT, down-projected on to the corona for an assumed velocity of  $800 \text{ km s}^{-1}$ . The triangle corresponds to the point of closest approach of the line of sight to the Sun (and therefore the region which makes the maximum contribution to the measured scintillation) and the circle corresponds to the Earth. The tick-marks indicate increments of  $10^\circ$  in  $\theta$  along the line of sight, showing that the signal from the source remains in the region above the polar coronal hole, past the point of closest approach to the Sun, until  $\theta$  equals  $-45^\circ$  (see Fig. 1). The recognition of coronal holes as the source regions for fast streams in the solar wind is “well-established” (Axford and McKenzie, 1996) so we would expect observations of 0432 + 416 centred on 11 June 1995 to indicate a high speed solar wind. Figure 9b shows the auto- and cross-correlation functions observed by EISCAT: the correlation functions were fitted after making a further improvement in the integration model which uses white-light information to

locate the position of the fast- and slow-speed streams along the line of sight as indicated by the parameter  $\theta$  (Grall *et al.*, 1996). The results from fitting this new model show a simple fast stream with a mean (true) velocity of  $813 \text{ km s}^{-1}$  and small rms variation in the transverse velocity  $V_t$  of  $14 \text{ km s}^{-1}$ . No attempt has been made to fit the slow solar wind as it is virtually impossible to extract information on the slow component when scattering is clearly dominated by the fast wind. These results are entirely consistent with the picture of a smooth fast solar wind flowing out from polar coronal holes which has emerged from Ulysses observations (Geiss *et al.*, 1995; Phillips *et al.*, 1995).

Figure 10a shows the white-light map based on east-limb HAO observations, again centred on 11 June 1995. In this case the line of sight for observations of 0603 + 203 are mapped back into the central streamer belt with an assumed velocity of  $400 \text{ km s}^{-1}$ . The observed correlation functions have been fitted with a line of sight integration model which assumes a single slow stream (Fig. 10b). The





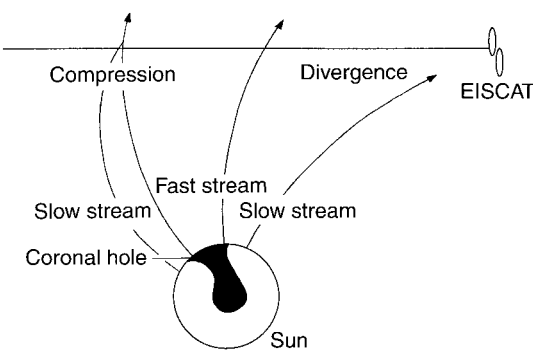
**Fig. 10. a** White-light map with the IPS line of sight for observations of 0603 + 219 on 11 June 1995 (a single slow stream) back-projected onto the map with an assumed constant radial velocity of  $400 \text{ km s}^{-1}$ . **b** IPS correlation function corresponding to **a**, analysed using the line of sight model. The *dashed line* represents the data and the *solid line* the model fit

model represents the observed correlation functions very well, suggesting a mean velocity of  $405 \text{ km s}^{-1}$  and an rms variation in  $V_t$  of  $99 \text{ km s}^{-1}$ .

These results are typical of the data analysed using the line of sight integration model. In all cases studied so far, simple fast streams are associated with coronal holes and simple slow streams with bright “non-hole” regions of the corona. These results provide a sound basis for approaching the more complex problem of interpreting observations where two streams are present along the line of sight.

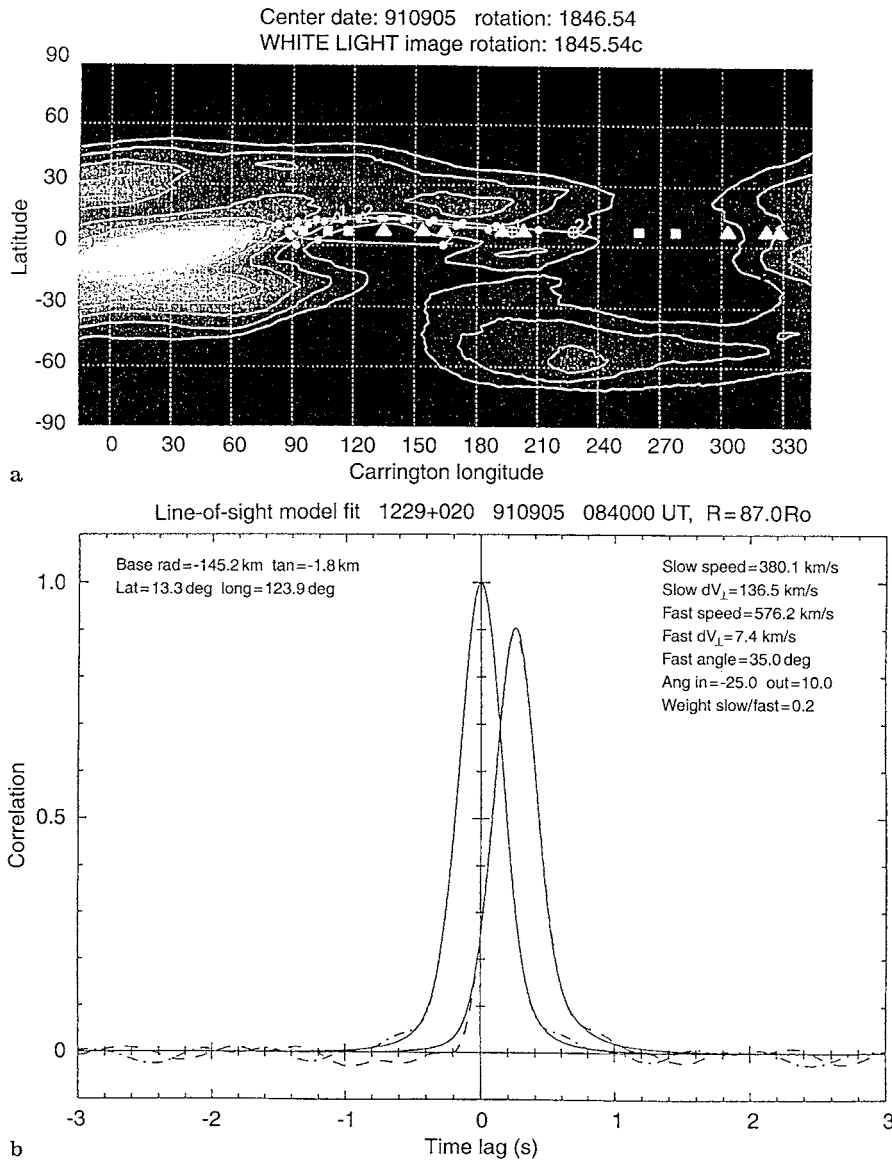
2.5 Fast stream/slow stream interaction regions

Although the plasma of the solar wind flows out radially, the source regions of the solar wind follow the rotation of the Sun. If there is a coronal hole extending to middle or low heliospheric latitudes, bordered by regions of closed magnetic field lines, the rotation of the Sun will carry the



**Fig. 11.** Schematic diagram of fast stream/slow-stream interaction, looking down on the north pole of the Sun

region of open flux underneath a region of outward-flowing slow wind (Fig. 11). Under these circumstances the fast stream will overtake the slow stream at one boundary, and diverge from it at the other.



**Fig. 12.** **a** White-light map for 5 September 1991 showing the line of sight for measurements of 1229 + 020 back-projected onto the map with assumed velocities of (1)  $800 \text{ km s}^{-1}$  and (2)  $400 \text{ km s}^{-1}$ . The fast stream is overtaking the slow stream; **b** IPS correlation function corresponding to **a**, analysed using the line of sight model. The fast stream is much slower and causes much more scintillation than expected, suggesting that the fast stream and the slow stream are interacting to produce a turbulent compression zone. The *dashed line* represents the data and the *solid line* the model fit

In the case where the fast stream overtakes the slow stream, compression will occur and the interaction region will become turbulent. It can also be predicted that the fast stream will decelerate and the velocity become non-radial, while the compressed turbulent region will produce strong scintillation (private communication, Wilkinson, 1995). On the other boundary of fast/slow interaction a rarefaction zone would be expected to develop: in this region both fast and slow streams will be in the line of sight but without any strong interaction and two undisturbed streams should be seen.

Figure 12a shows the white-light map of the west limb of the Sun centred on 5 September 1991. The superimposed lines show the lines of sight from EISCAT to the radio source 1229 + 020 projected onto the corona with assumed velocities of (1)  $800 \text{ km s}^{-1}$  and (2)  $400 \text{ km s}^{-1}$ . Both projections take the line of sight across the boundaries of a low-latitude coronal hole, with the sector above

the coronal hole lying approximately between  $\theta = +10^\circ$  and  $\theta = -25^\circ$ . The fast stream is overtaking the slow stream on one boundary of this hole. The corresponding EISCAT observations and the line of sight model fit are shown in Fig. 12b. The fast component is anomalously slow at  $576 \text{ km s}^{-1}$  and is weighted more heavily than the slow stream, which is contrary to the usual observations of undisturbed fast and slow streams (Grall *et al.*, 1996) but consistent with enhanced levels of scintillation within a turbulent, compressed, interaction region.

The slightly off-radial velocity observed on 5 June 1995 (Fig. 8c) also appears to be associated with a fast stream overtaking a slow stream. The closest approach of the line of sight to the Sun lies just inside the polar coronal hole and fitting with the line of sight integration model gives good agreement with the EISCAT observations. The model results show no clear evidence of compressional interaction between the fast and slow streams, but in

this case the line of sight is considerably closer to the Sun (closest approach =  $32 R_{\odot}$ ) and so the interaction between the fast and slow streams might be expected to be much weaker.

When the fast and slow streams of solar wind are diverging from each other a rarefaction region is produced. In this region the low density of plasma (and the reduced sharpness of the density gradients) reduces the levels of scintillation. Rarefaction regions cannot, therefore, be directly detected by IPS measurements, but their presence can be inferred from observations which show undisturbed fast and slow components of the solar wind coming from regions where the flow from a coronal hole appears to be diverging from the solar wind.

### 3 Discussion and conclusions

Earlier studies of IPS observations assumed that the solar wind contained a range of velocities between about 400 and  $800 \text{ km s}^{-1}$ , but more recent spacecraft observations have suggested that the solar wind close to the Sun is best described as consisting of distance high and low speed modes (Schwenn, 1990). EISCAT allows the fast and slow streams to be resolved, something which is not possible with single-site observations or with two-site observations on shorter baselines (Grall *et al.*, 1996), and the smooth range of velocities reported in earlier work resulted from differing proportions of fast and slow streams in the line of sight through the solar atmosphere. The observations presented in this paper and by Grall *et al.* (1996) confirm that the solar wind is bimodal with clear fast ( $\sim 800 \text{ km s}^{-1}$ ) and slow ( $\sim 400 \text{ km s}^{-1}$ ) components. The slow stream dominates the data at low heliographic latitudes, while only fast streams are seen at high latitudes. These results are in complete agreement with Ulysses observations (Phillips *et al.*, 1994, 1995). When corrected for foreshortening the velocities observed by EISCAT agree well with the fast- and slow-stream velocities seen by Ulysses. Down-projection from the IPS line of sight on to white-light maps of the corona demonstrated the relationship between stream characteristics and coronal features. As expected there is a clear association of fast streams with areas of open magnetic flux (coronal holes), while slow streams are associated with the “non-hole” areas of the corona and the small proportion of intermediate velocities can be interpreted as the result of fast-stream/slow-stream interaction. This interpretation is supported by observations of a fast stream overtaking a slow stream in which the fast stream appears to have decelerated and the level of scintillation increased.

EISCAT measurements over a wide range of distances from the Sun ( $15\text{--}130 R_{\odot}$ ) show that the velocity of the irregularities producing the scintillations does not vary systematically with distance: if it is assumed that these irregularities are carried at the bulk flow speed of the solar wind plasma then the solar wind has completed its acceleration inside  $15 R_{\odot}$  for fast streams or inside  $25 R_{\odot}$  for slow streams. Indeed a series of co-ordinated observations using EISCAT and the VLBA (Grall *et al.*, 1996) showed that the fast high-latitude solar wind reaches its final

velocity inside  $10 R_{\odot}$ , while velocities inferred from SPARTAN-201 observations suggest that the acceleration region may lie inside  $5 R_{\odot}$ . These results place significant constraints on models seeking to explain the acceleration of the solar wind (Grall *et al.*, 1996), although a simple model (McWhirter and Kopp, 1979) predicts that the main acceleration region should be inside  $5 R_{\odot}$ .

It remains possible that IPS observations are sensitive to waves in the solar wind, with the observed scintillation being produced by the interactions of Alfvén waves with the anisotropic medium. In this case the density fluctuations producing the scintillation would not be travelling at the bulk flow speed. If this is the case then the velocity of the ‘scintillators’ remains constant over a range of distances covering  $10\text{--}130 R_{\odot}$  and is comparable to the bulk velocity of the plasma observed by Helios at  $60 R_{\odot}$  and at up to  $600 R_{\odot}$  by Ulysses. Grall *et al.* (1996) attempted to estimate the bias to the velocity measurements introduced by Alfvén waves travelling through an anisotropic microstructure and concluded that this was not a serious problem for measurements outside  $6 R_{\odot}$ . This suggests that IPS is capable of measuring the bulk velocity of the solar wind at distances outside  $10 R_{\odot}$ , the closest observation to the Sun discussed in this study.

Regions of interaction between fast and slow streams have been observed. Preliminary studies suggest that compressional interactions, produced when a fast stream overtakes a slow stream, might be the origin of the small proportion of velocities observed which fall between the normal fast- and slow-stream speeds. The flow of the solar wind may also become slightly off-radial in compressional interaction regions.

The recent series of IPS observations confirm EISCAT as an extremely powerful instrument for studying the solar wind. Normally the interpretation of effects that have been integrated along a line of sight is fraught with difficulties, but the bimodal distribution of solar wind velocity and the clear association of fast streams with coronal holes allow the analysis programme to integrate along the line of sight, crossing between fast-stream and slow-stream regimes at boundaries whose position can be derived from white-light coronagraph measurements. The success of this analysis programme in fitting the observed data to provide physically believable and consistent parameters fully justifies our confidence in the results obtained. It is now intended to use the same combination of EISCAT data and white-light maps to analyse the entire 1990–1995 dataset as part of a survey of the three-dimensional character of the solar wind. At the same time the successful launch of SOHO and planned improvements in the performance of the EISCAT system promise even better results in the future.

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